



Slow neutrons and their role to test gravity

Testing Gravity 2015

16.01.2015

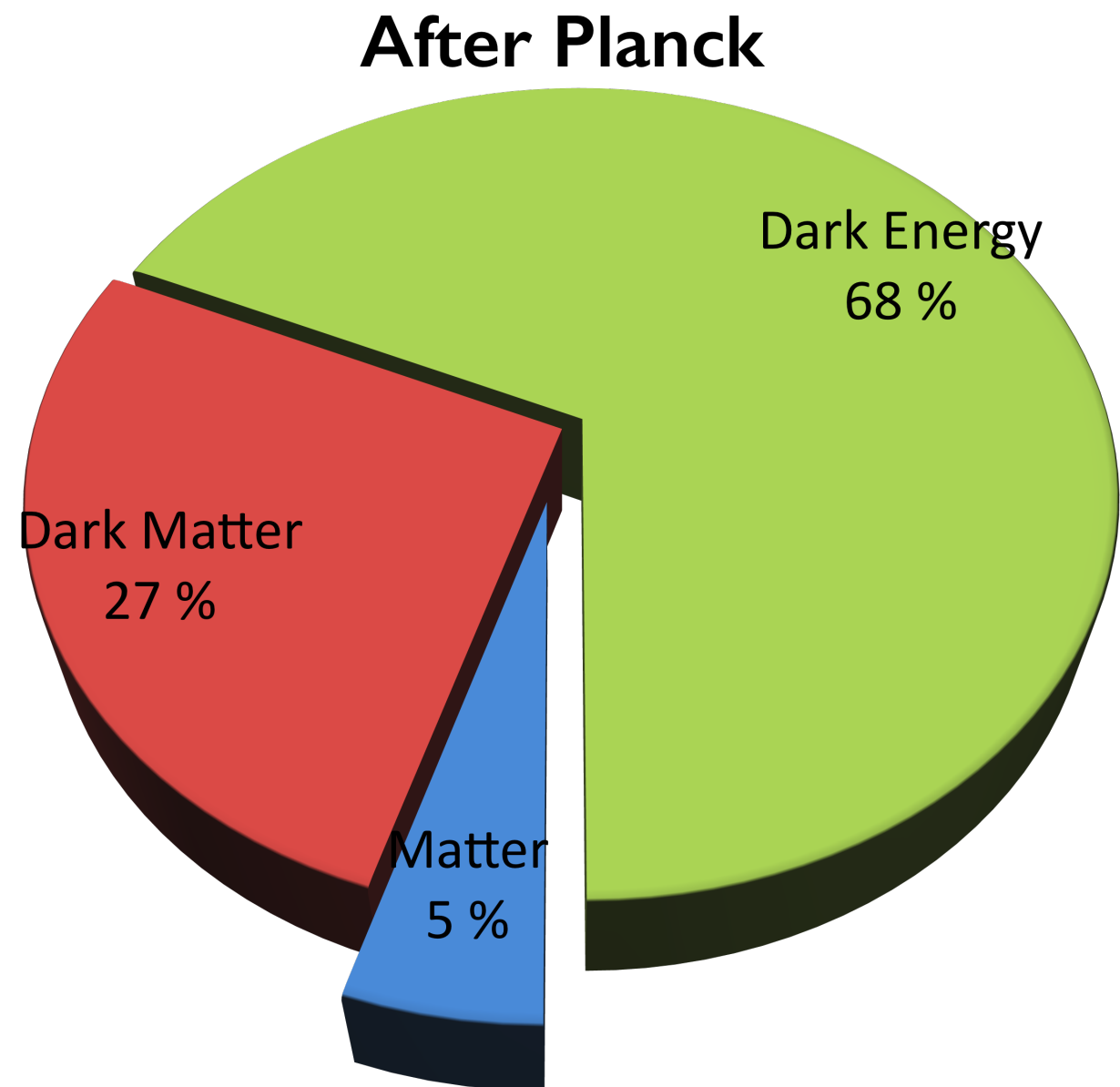
Gunther Cronenberg



Dark Matter & Dark Energy Search

- galaxy rotation problem
- Supernova Ia
- cosmic microwave background (CMB)

experimental evidence for
Dark Matter (DM) and
Dark Energy (DE)



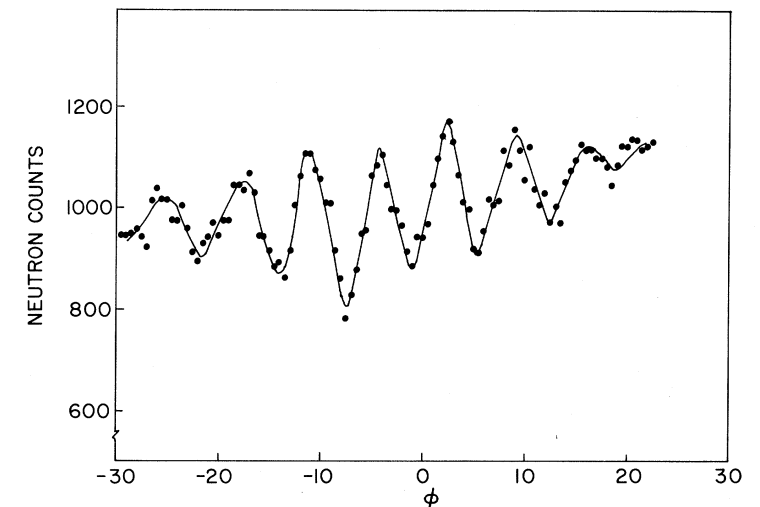
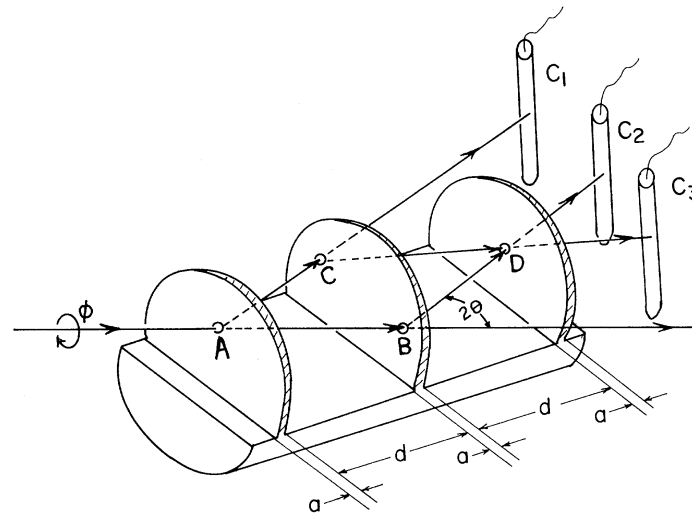
Neutrons & Gravity



First experiments:

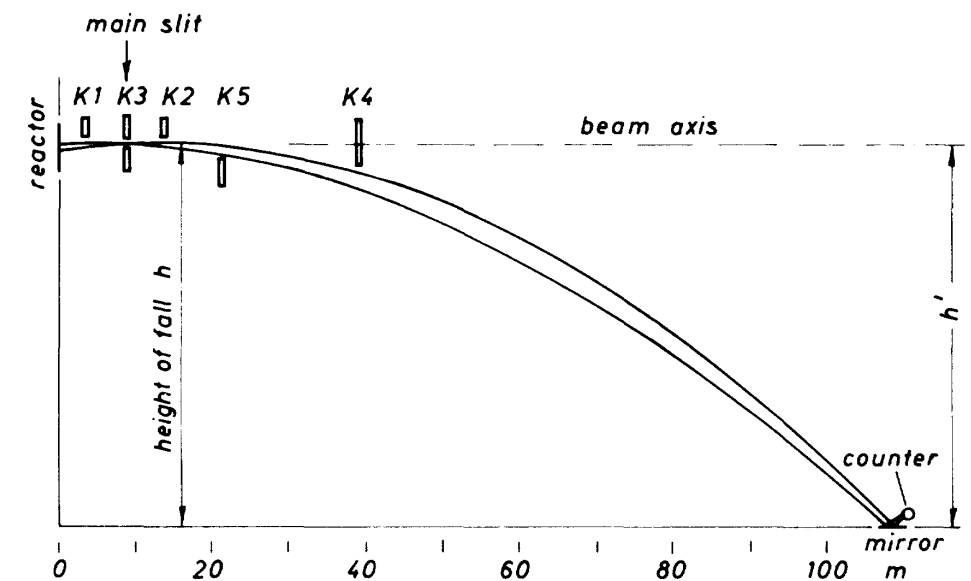
- COW (1975):

first observation of
gravitational shift
for neutrons



- Köster (1976):

test equality of inertial and
gravitational mass
by scattering length measurements



Colella, R., Overhauser, A., & Werner, S. (1975). Physical Review Letters, 34(23), 1472

Koester, L. (1976). Physical Review D, 14(4),

Ultra/Very Cold Neutrons

UCN $\lambda \approx 1.5\text{nm} - 15\text{nm}$
VCN $\lambda \sim 50\text{nm}$ $E_{\text{kin}} \approx (50 - 300)\text{neV}$

Neutron:

- no electric charge
- small polarisability
- nEDM

$$(11.6 \pm 1.5) \cdot 10^{-4} \text{ fm}^3$$

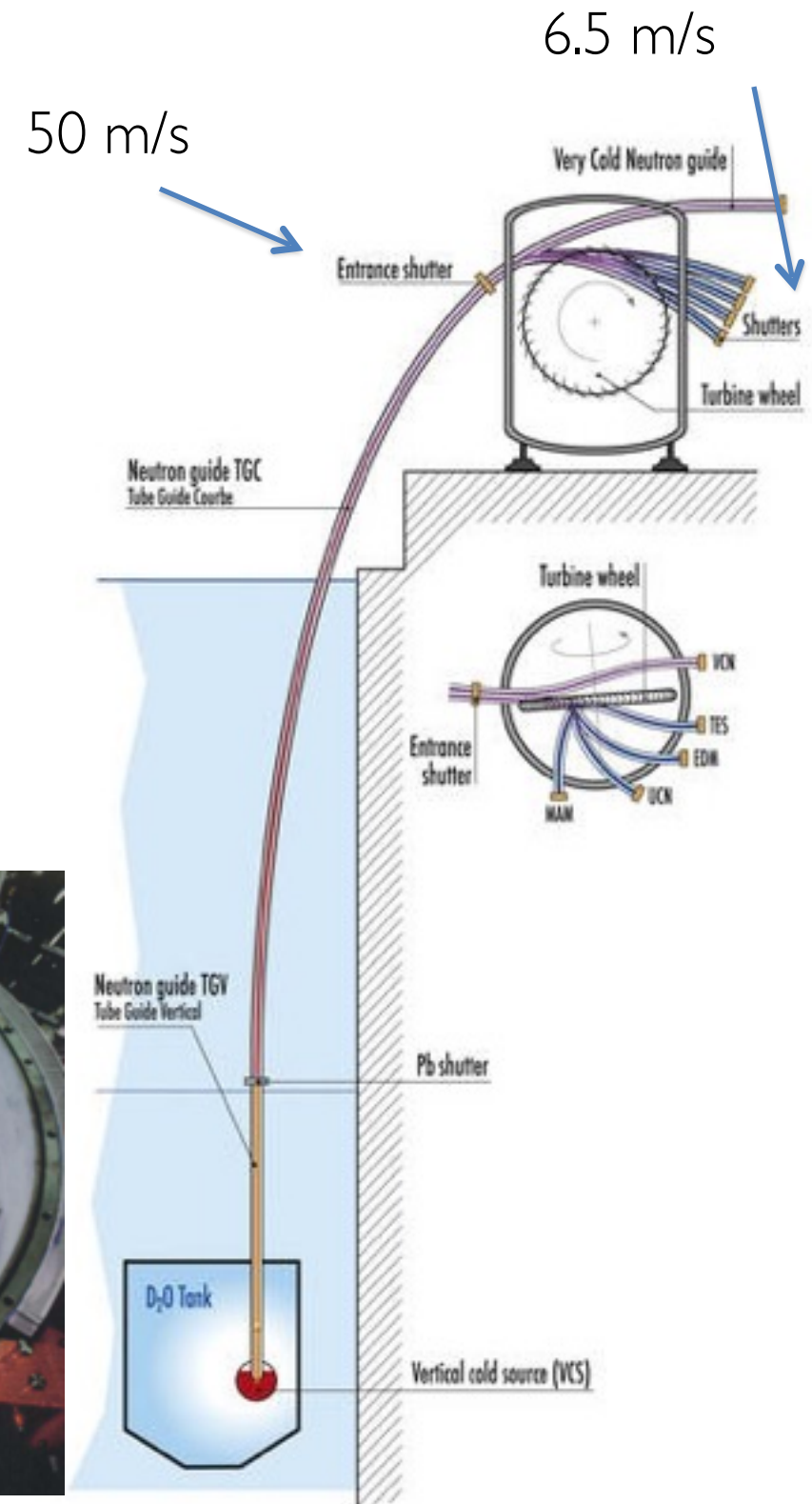
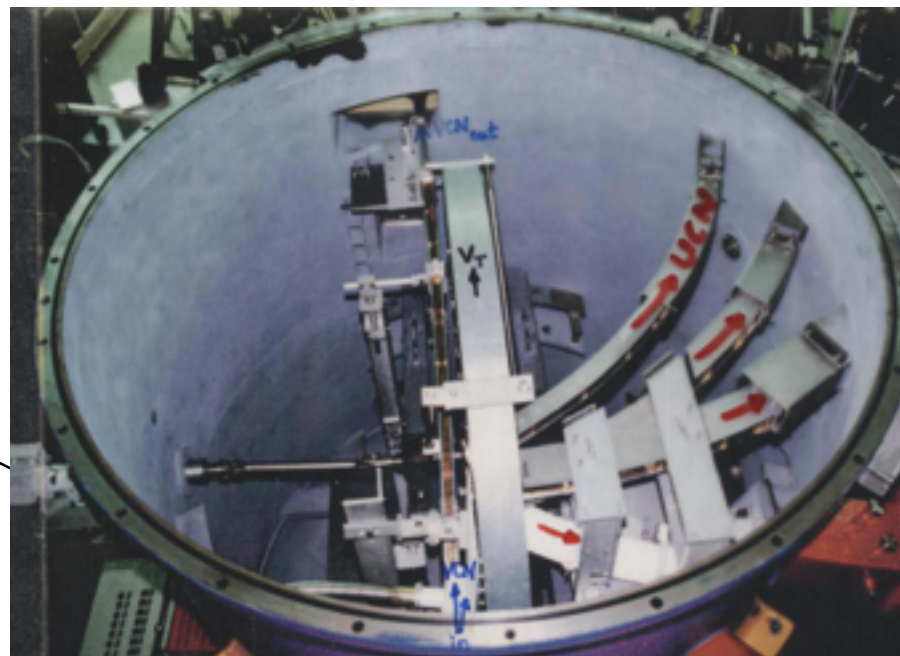
$$\propto 10^{-19} \alpha_{\text{atom}}$$

$$|d_n| < 2.9 \times 10^{-26} \text{ ecm}$$

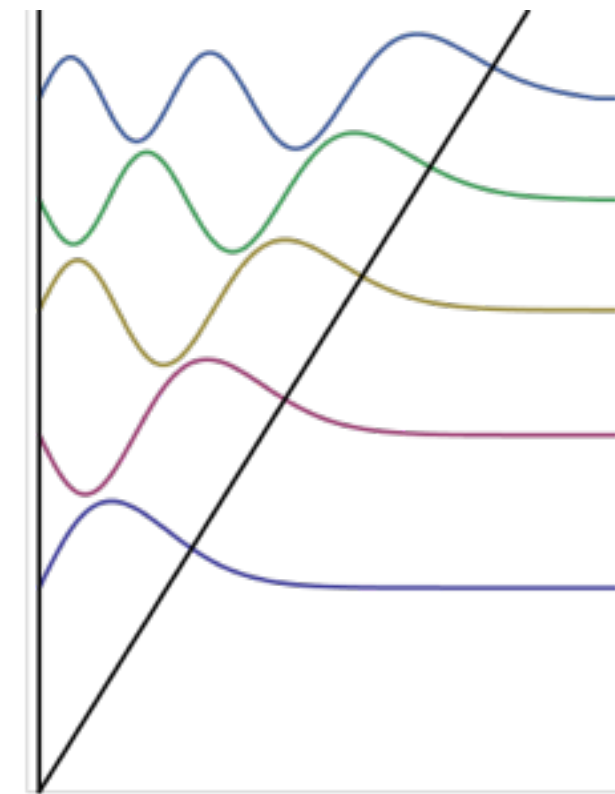
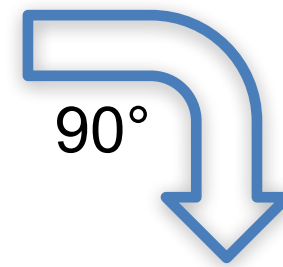
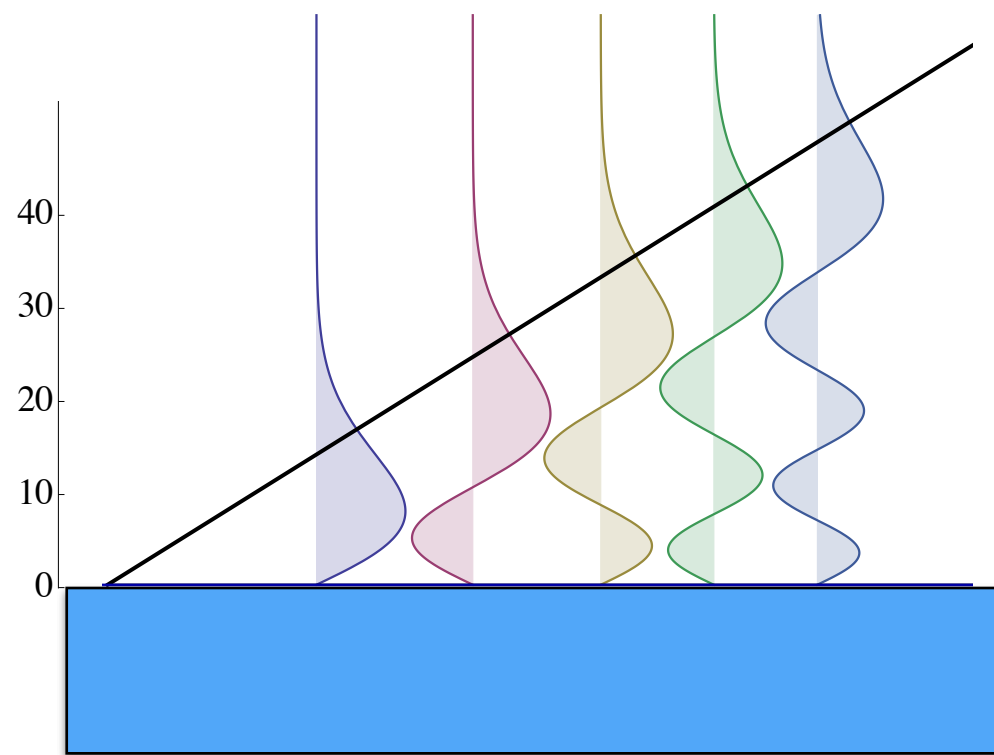
- total reflection for even big angles (VCN $< 1^\circ$)
- absorption !

UCN/VCN source:
PF2 at ILL

courtesy of C. Plonka



UCNs in the gravity field



V.I. Luschikov and A.I. Frank, JETP Lett. 28 559 (1978)
V. Nesvizhevsky et al., Nature, 415 297 (2002)

UCNs in the gravity field



- Schrödinger eq. with linearized gravity potential

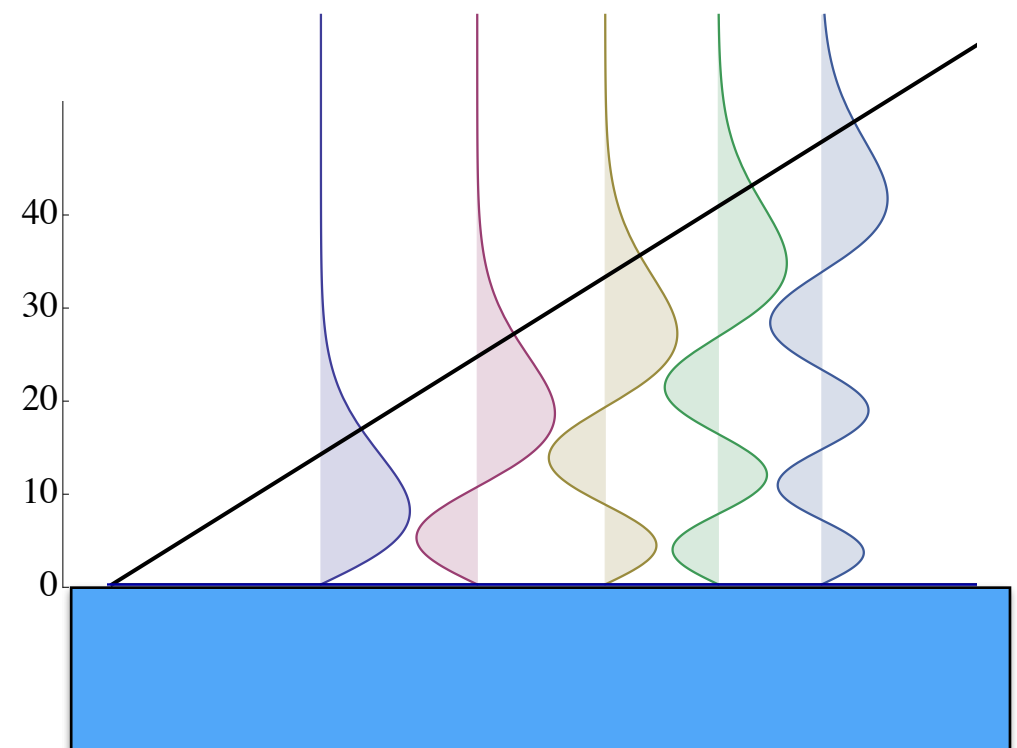
$$\left(-\frac{\hbar^2}{2m} \frac{\partial^2}{\partial z^2} + mgz \right) \varphi_n(z) = E_n \varphi_n(z)$$

$$\text{bc: } \varphi_n(0) = 0$$

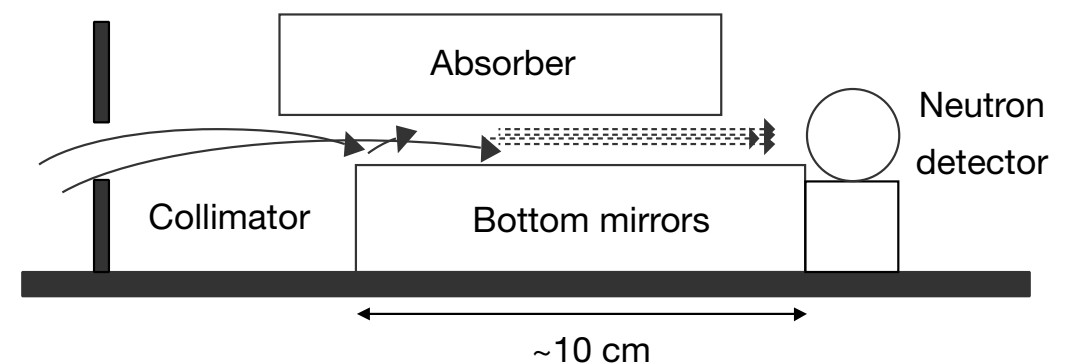
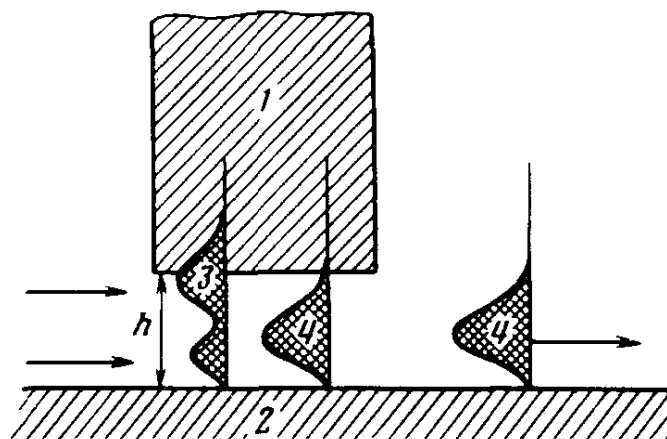
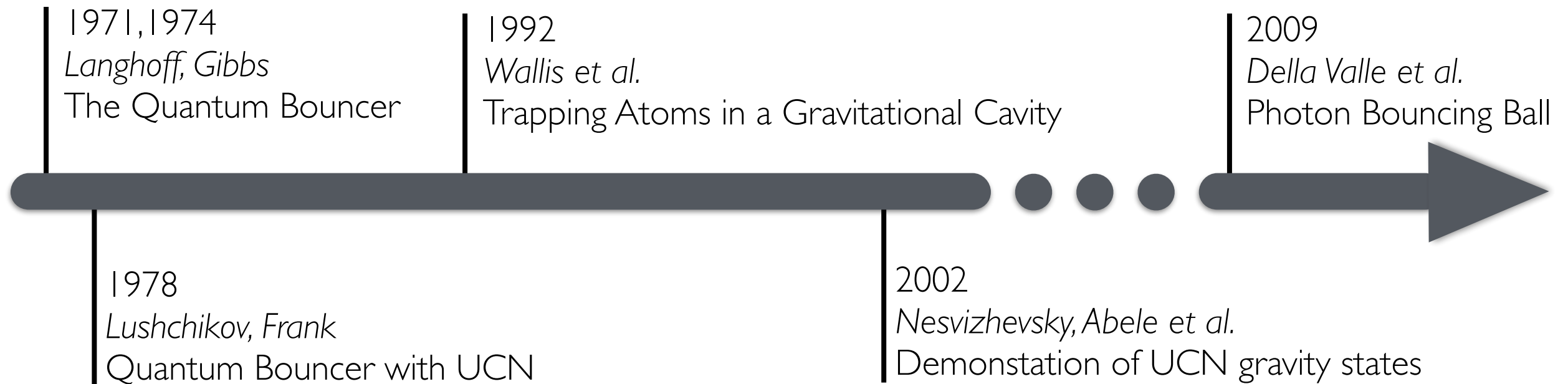
$$\varphi_n(z) = a_n \text{Ai} \left(\frac{z}{z_0} - \frac{E_n}{E_0} \right)$$

- bound, discrete states
- Non-equidistant energy levels

state	energy
1	1.41 peV
2	2.46 peV
3	3.32 peV



Quantum Bouncer



P.W. Langhoff, Am. J. Phys. 39, 954 (1971)
R. Gibbs, American Journ. o. Phys. 43 25 (1975)
V.I. Luschikov and A.I. Frank, JETP Lett. 28 559 (1978)

H.Wallis et al., Appl. Phys. B, 54 407 (1992)
V. Nesvizhevsky et al., Nature, 415 297 (2002)
G. Della Valle et al., PRL, 102 (2009)

qBounce: Timeline



2009

Jenke et al.

First Gravity Resonance Spectroscopy

2012

Cronenberg, TJ, HF, MT et al.

Full 3-part Rabi like setup

2014

Thalhammer, TJ, TR et al.

qBouncer

2007-2009

Jenke et al.

qBouncer

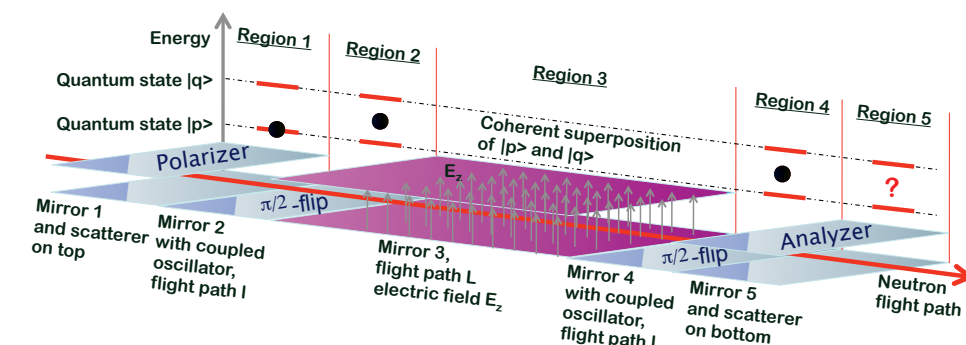
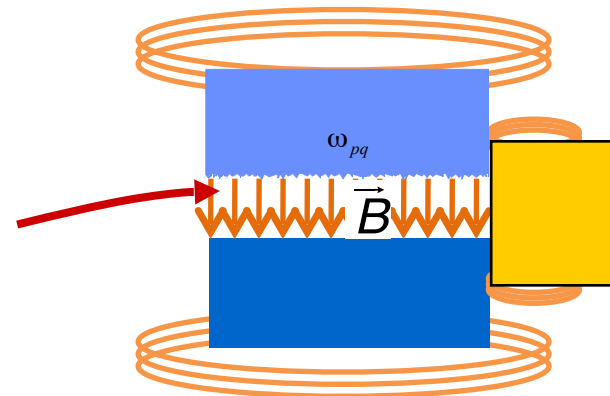
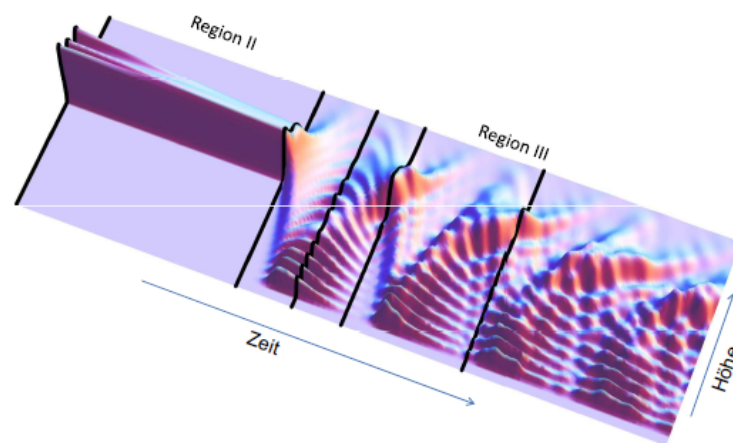
2010, 2011

Jenke, GC et al.

Axion, Chameleon measurements

H. Filter

Neutron charge



T. Jenke et al., NIM A 611 318 (2009)

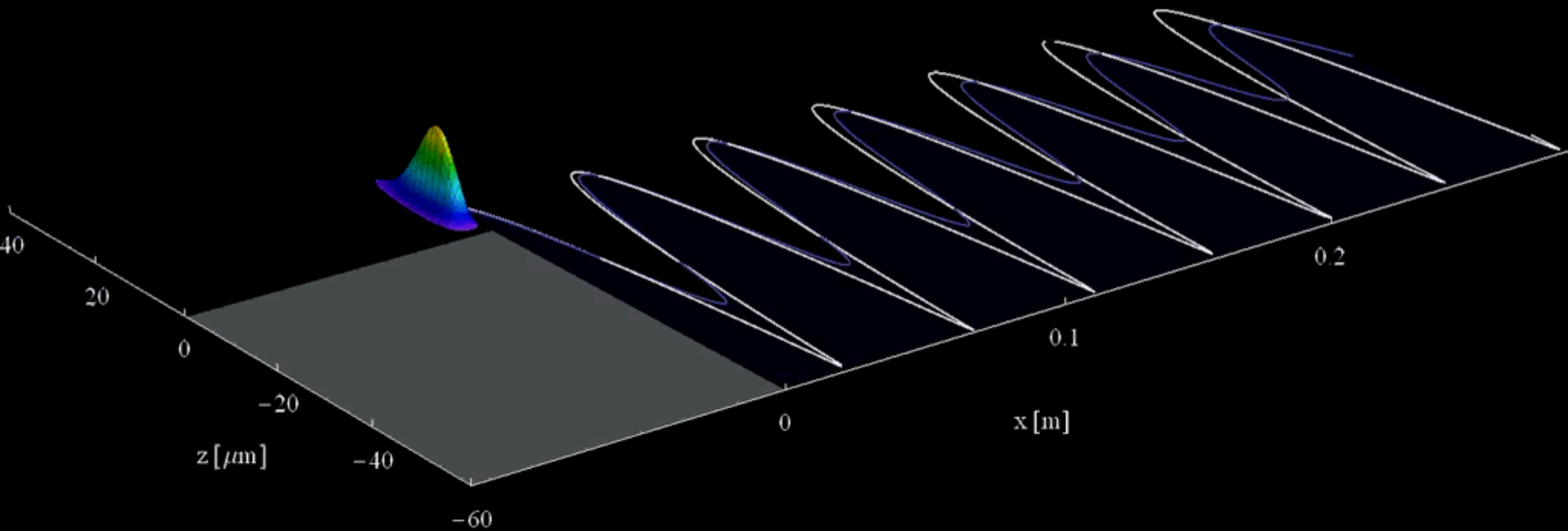
H. Abele et. al., Nucl. Phys A827, 593c (2009)

V.I. Luschikov and A.I. Frank, JETP Lett. 28 559 (1978)

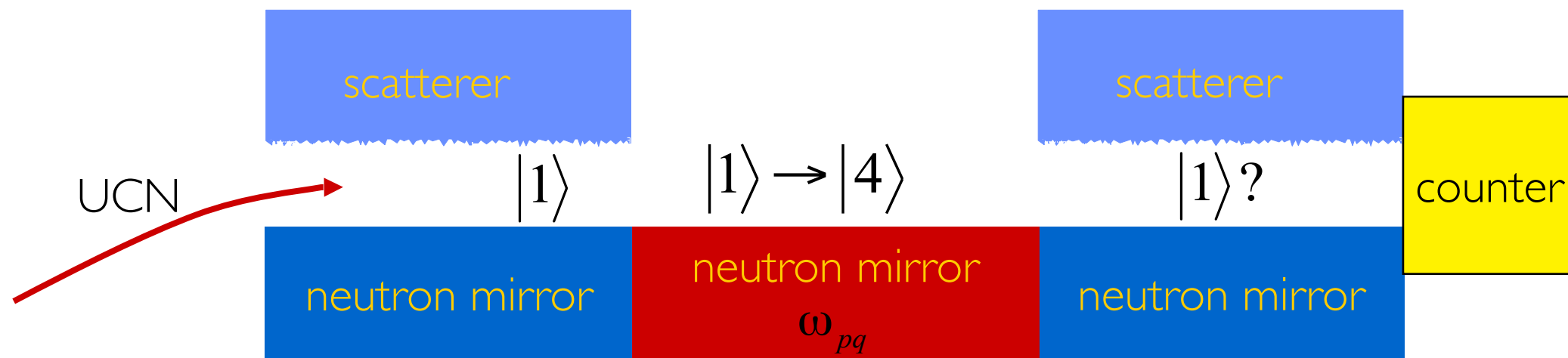
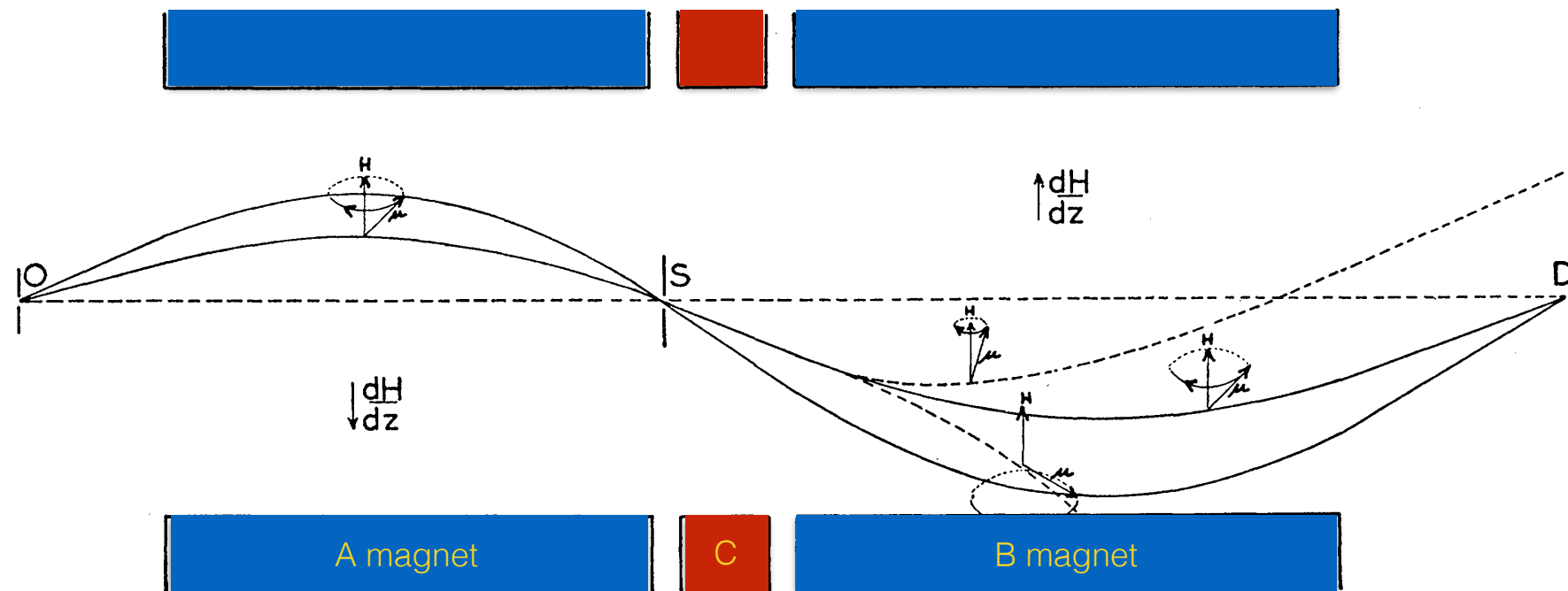
Quantum Bouncer

step: 60 μm

–1.000 ms



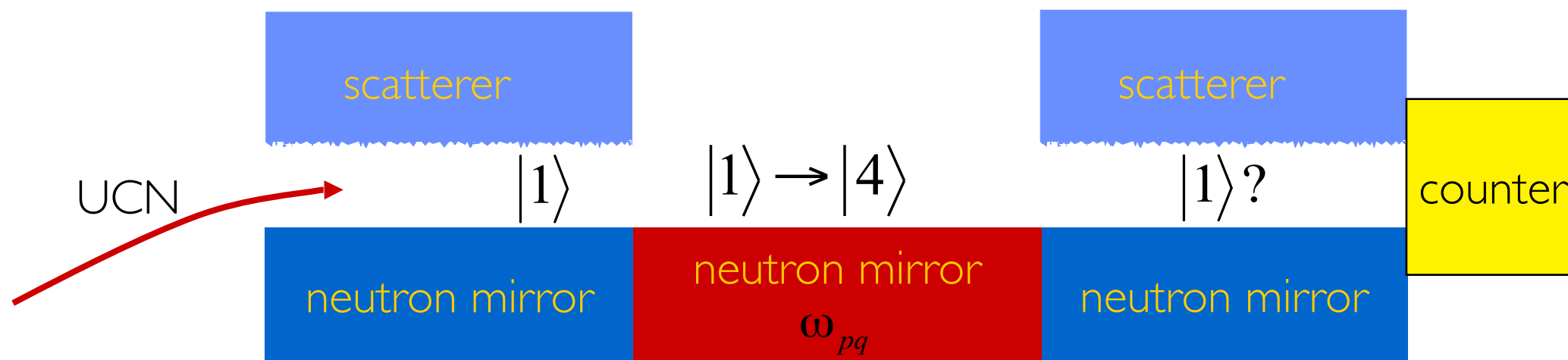
Rabis method



Gravity Resonance Spectroscopy

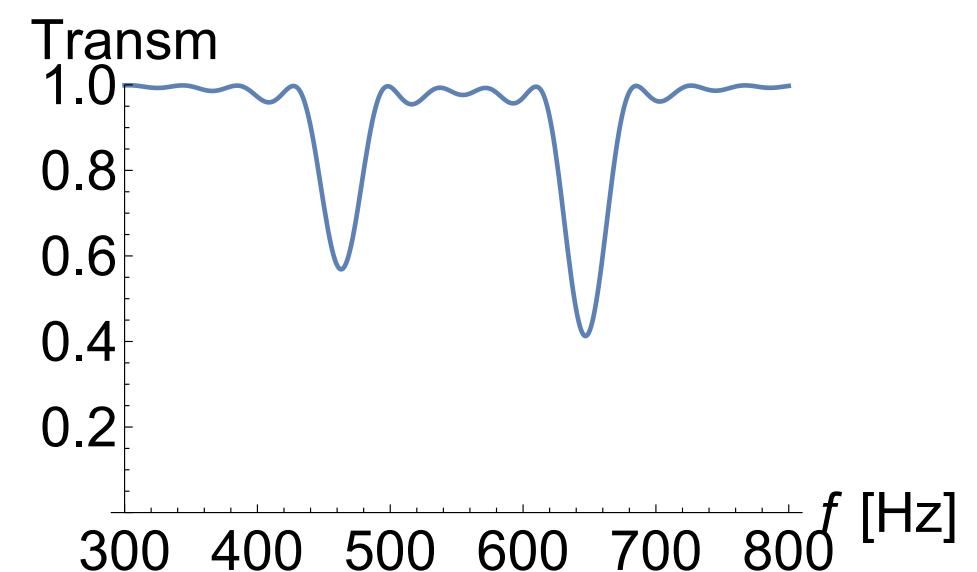
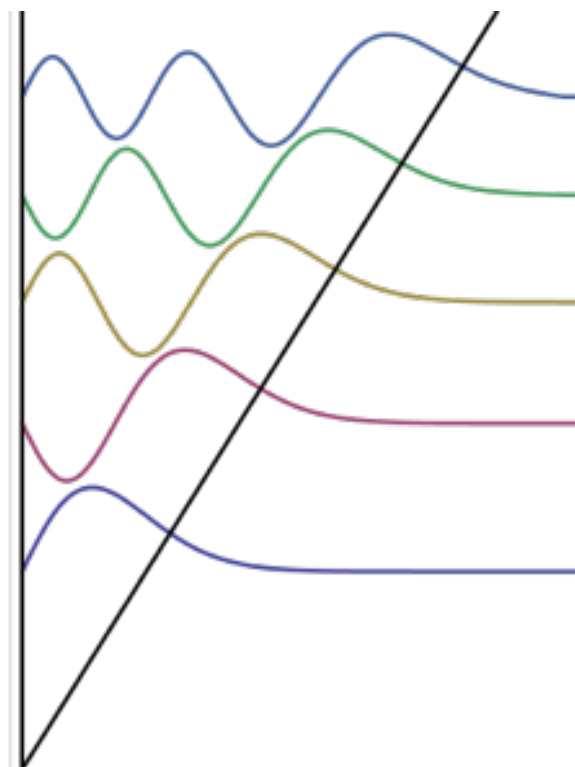


- Rabi



$$E = h\nu$$

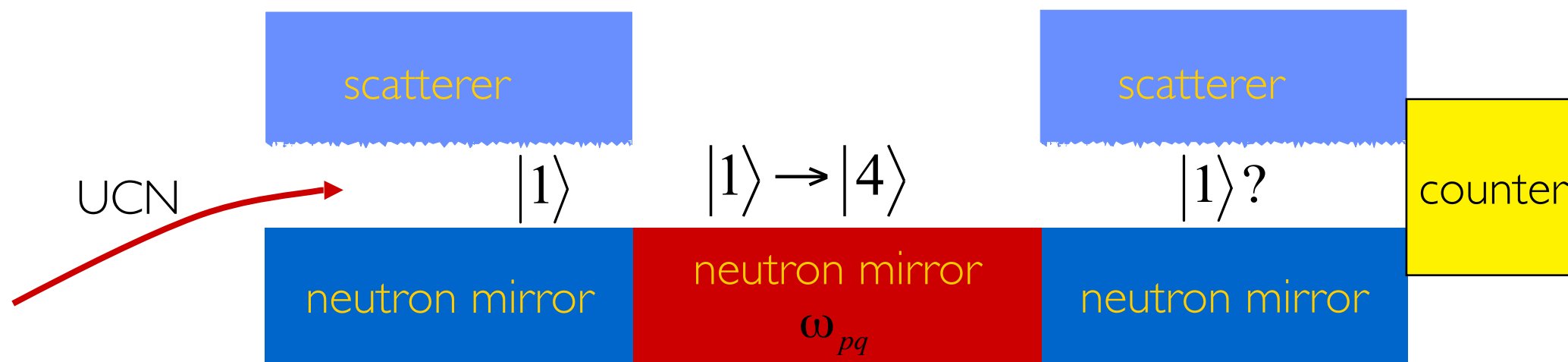
transitions induced by
mechanical oscillations or
magnetic fields



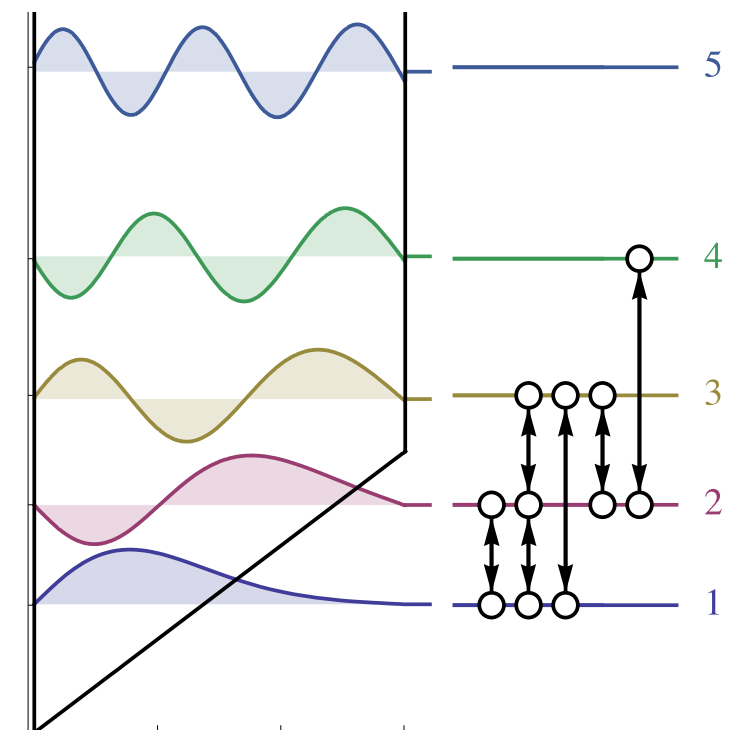
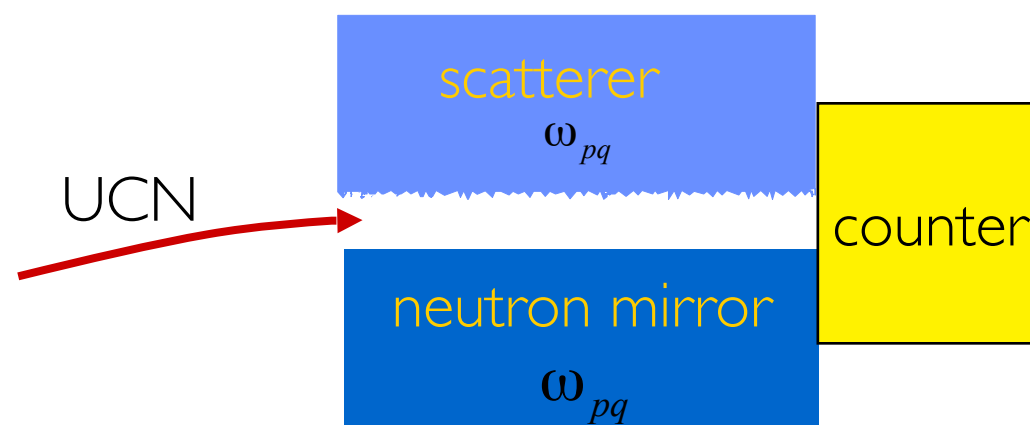
Gravity Resonance Spectroscopy



- Rabi (2012)



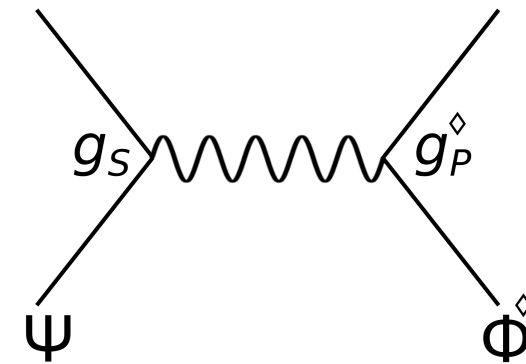
- First realisation (2009, 2010)



Axions



- spin-mass coupling
- scalar-pseudoscalar coupling



Introduced to solve problem on CP-violation in strong interactions

Candidates for dark matter

$$V(\vec{r}) = \hbar g_s g_p \frac{\vec{\sigma} \cdot \vec{n}}{8\pi m c} \left(\frac{1}{\lambda r} + \frac{1}{r^2} \right) e^{-r/\lambda}$$

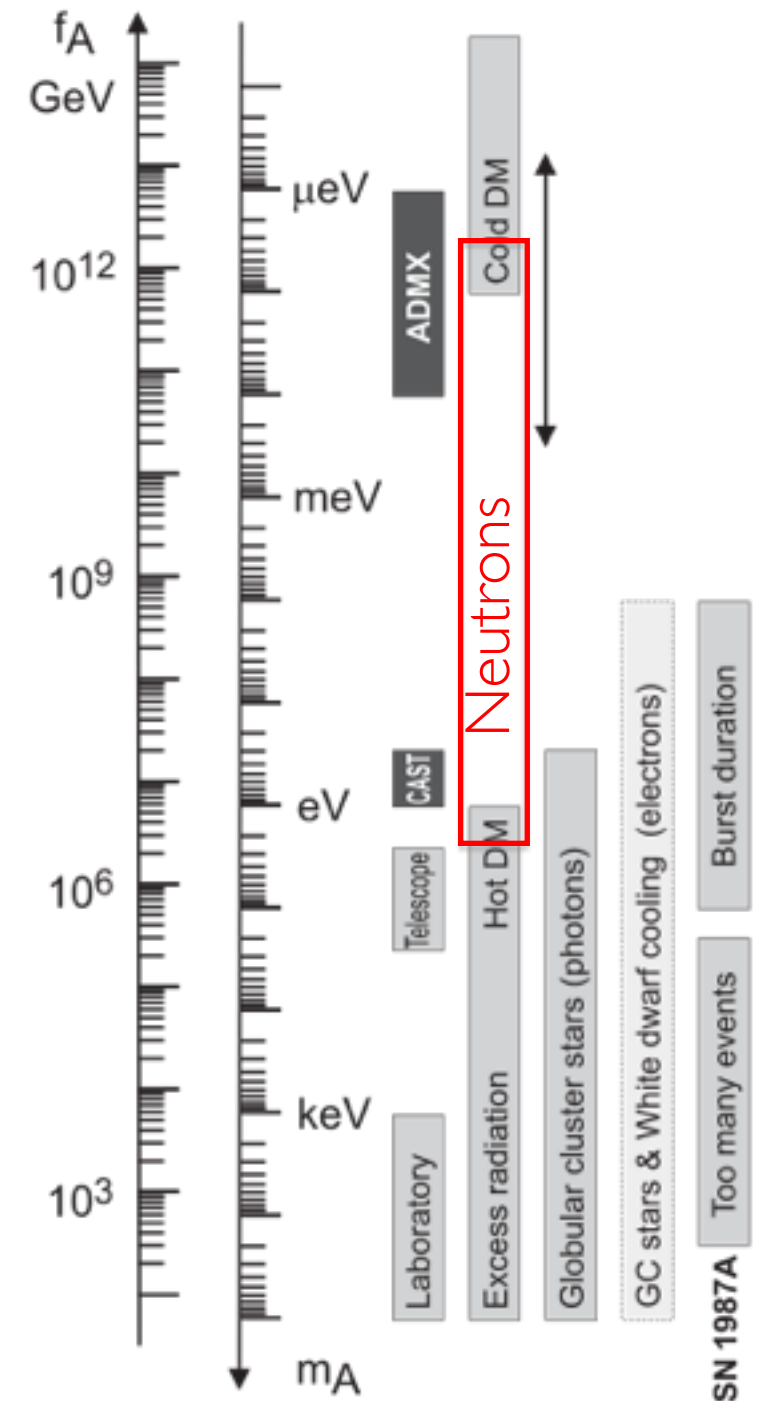
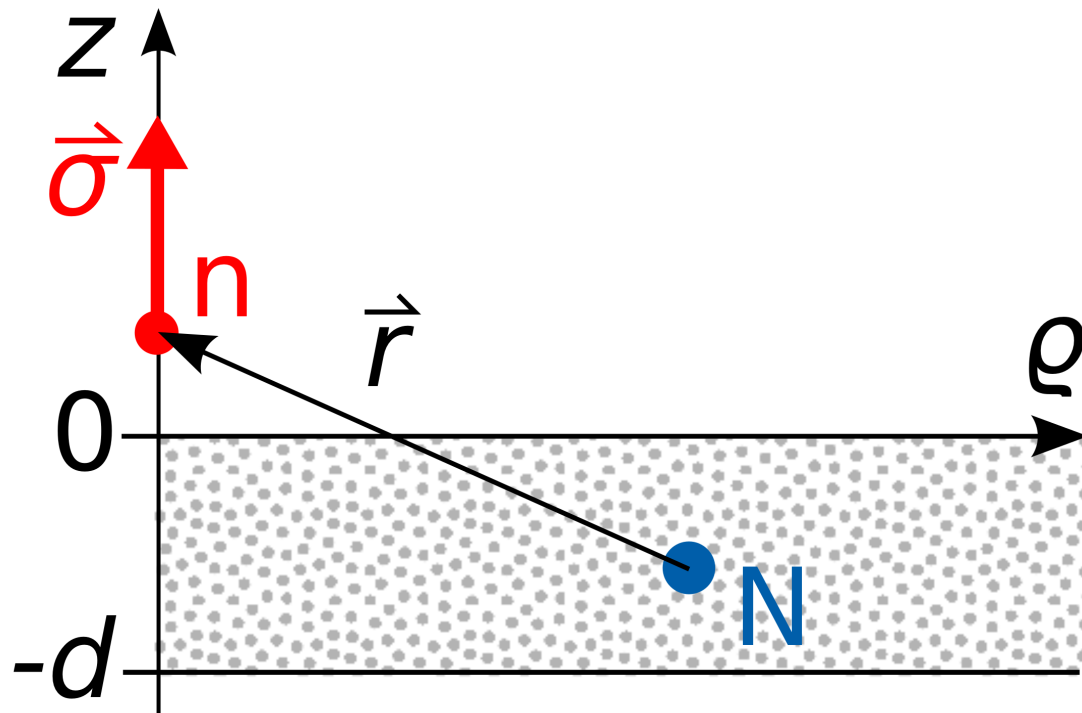
Also search for Axion-like particles (ALPs)

Search for Axions



Setup is sensitive to axions for

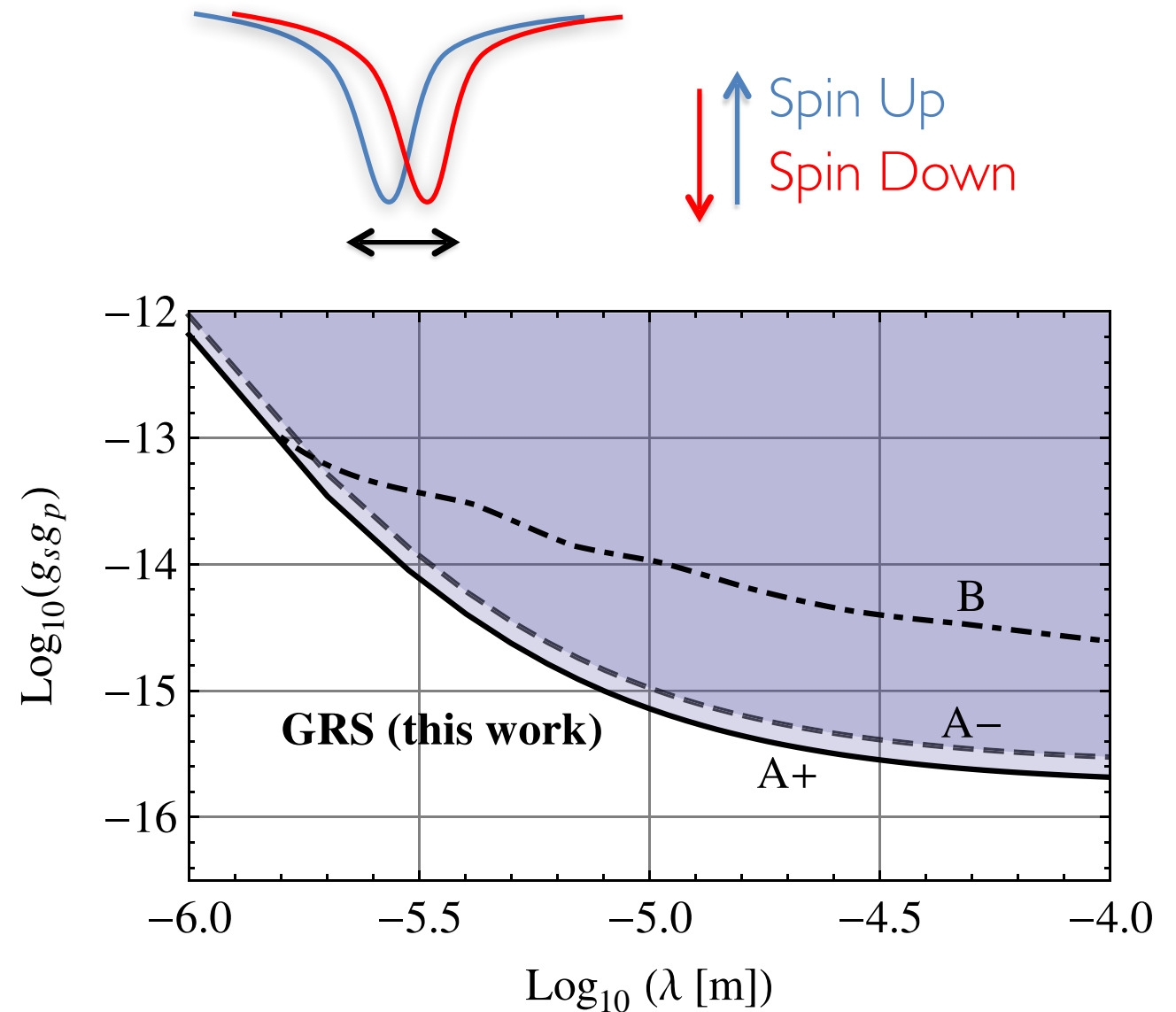
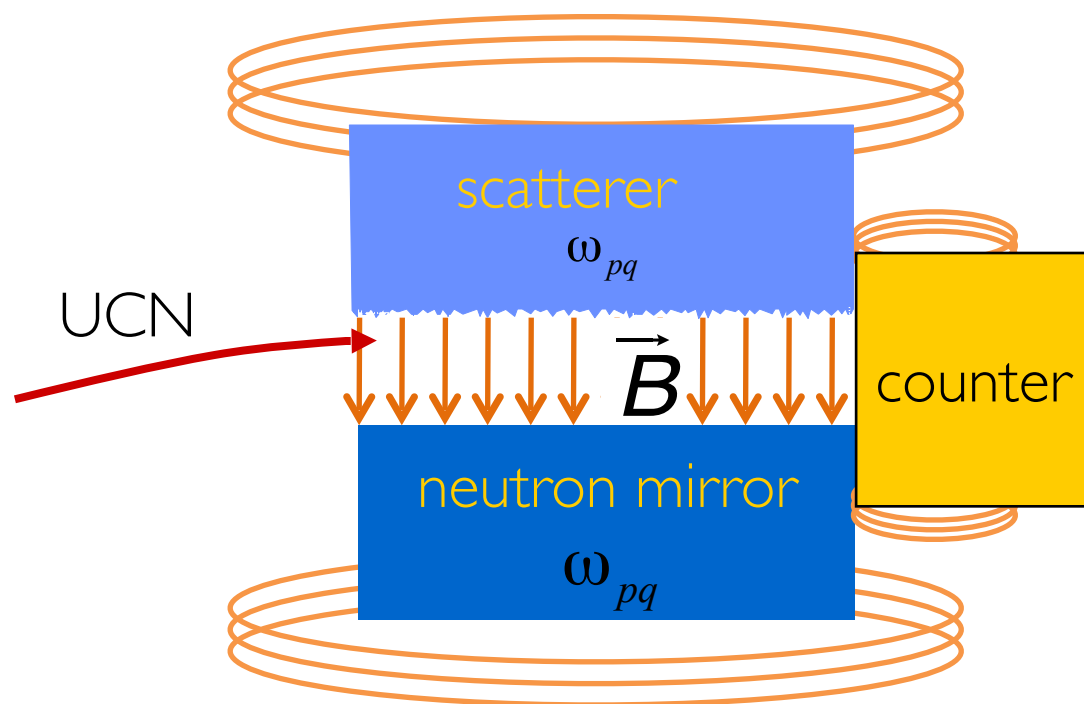
$$0.2\mu\text{m} \leq \lambda \leq 2\text{cm} \quad (10\mu\text{eV}..1\text{eV})$$



Search for Axions & ALPs



- Applying magnetic field
- Measuring effect for each spin





L. Filter



Chameleons



- Dark energy candidate
- mechanism suppresses in vicinity of masses

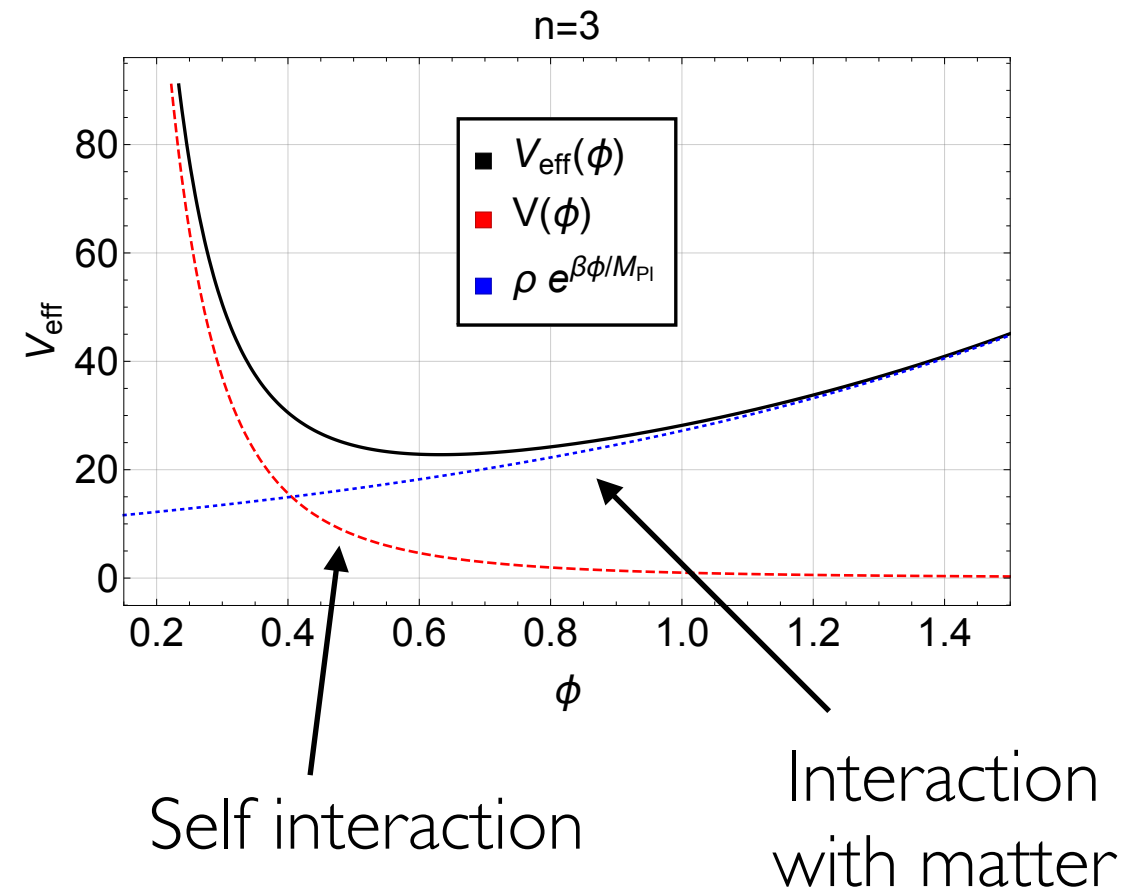
massless scalar field, hidden at laboratory scales

$$V_{\text{eff}}(\phi) = V(\phi) + \frac{\beta}{M_{\text{Pl}}} \phi \rho$$

Ratra-Peebles model ($n > 0$):

$$V(\phi) = \Lambda^4 + \frac{\Lambda^{4+n}}{\phi^n}$$

leads to screening effects





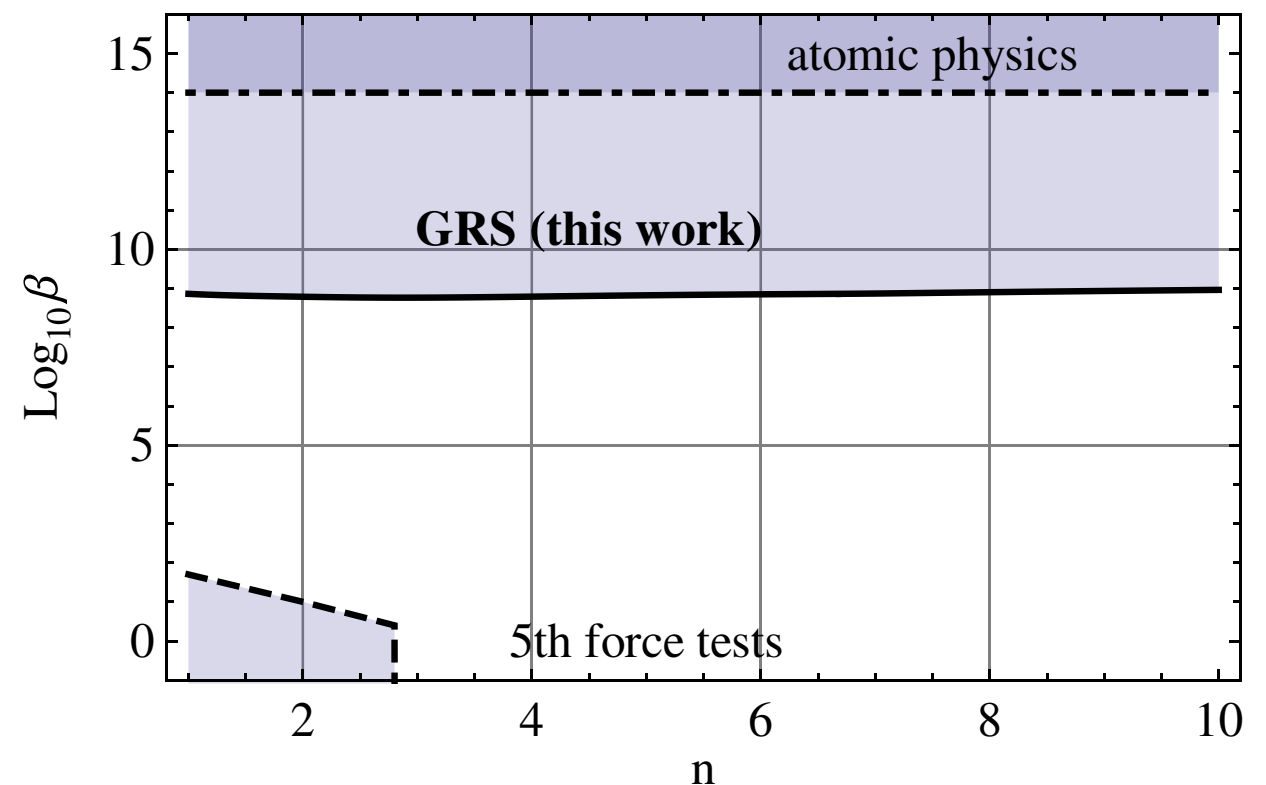
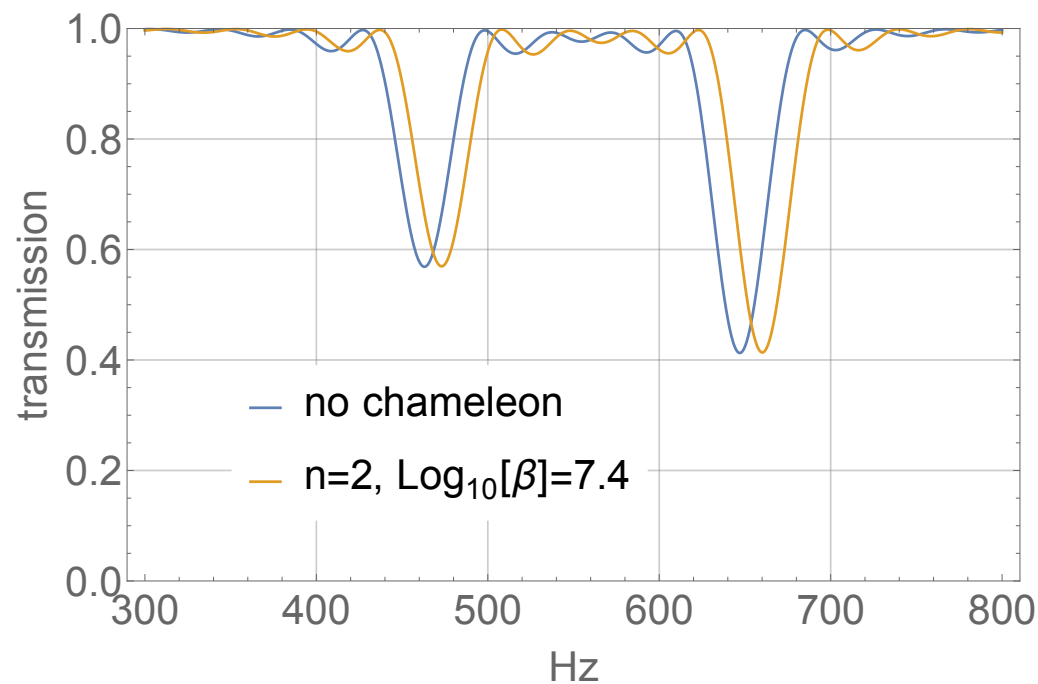
Chameleon fields



Chameleon shifts energy of state:

$$\delta E_n = \langle \psi_n | (z\Lambda)^{\frac{2}{2+n}} | \psi_n \rangle$$

$$V(z) = mgz + \beta \frac{m}{\overline{M}_{\text{Pl}}} \Lambda \hbar c \left(\frac{2+n}{\sqrt{2}} z\Lambda \right)^{\frac{2}{2+n}}$$



P. Brax, G. Pignol, Phys. Rev. Lett. 107, 111301 (2011)

A.N. Ivanov et al., (2013). Physical Review D, 87(10), 105013.

Jenke, T., Cronenberg, G., et al. Phys. Rev. Lett., 112, 151105. (2014)

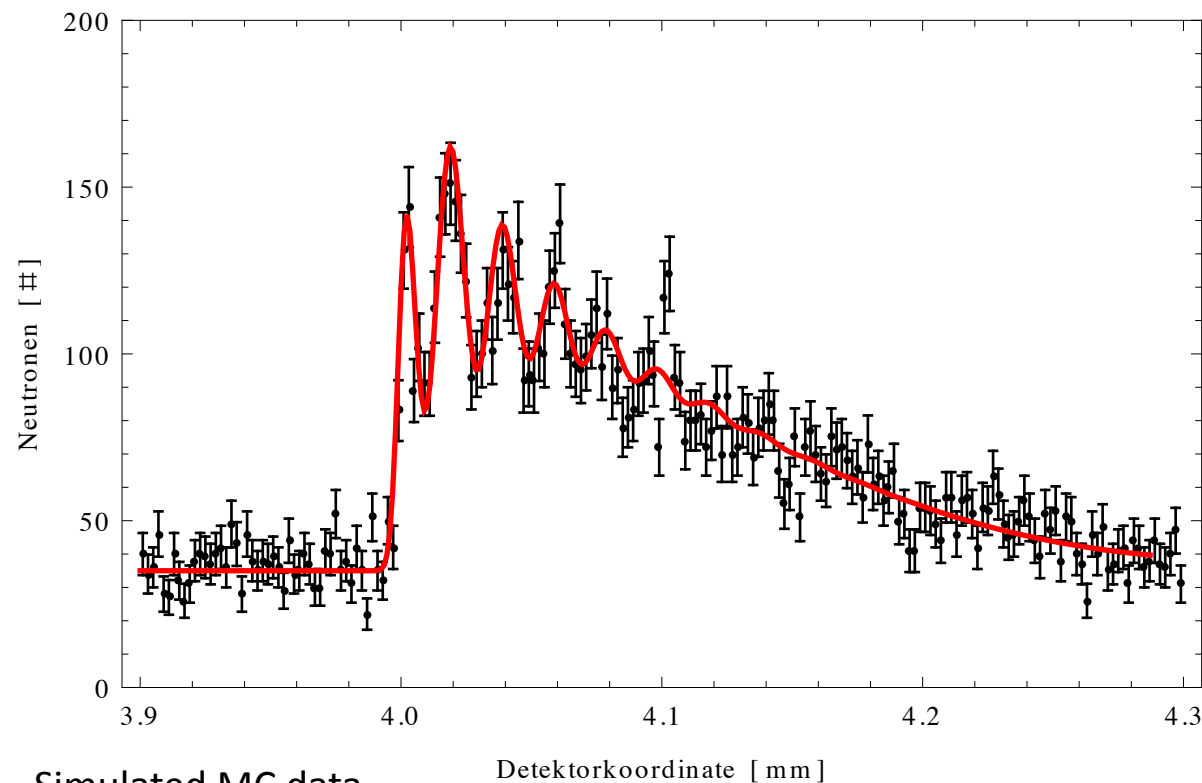
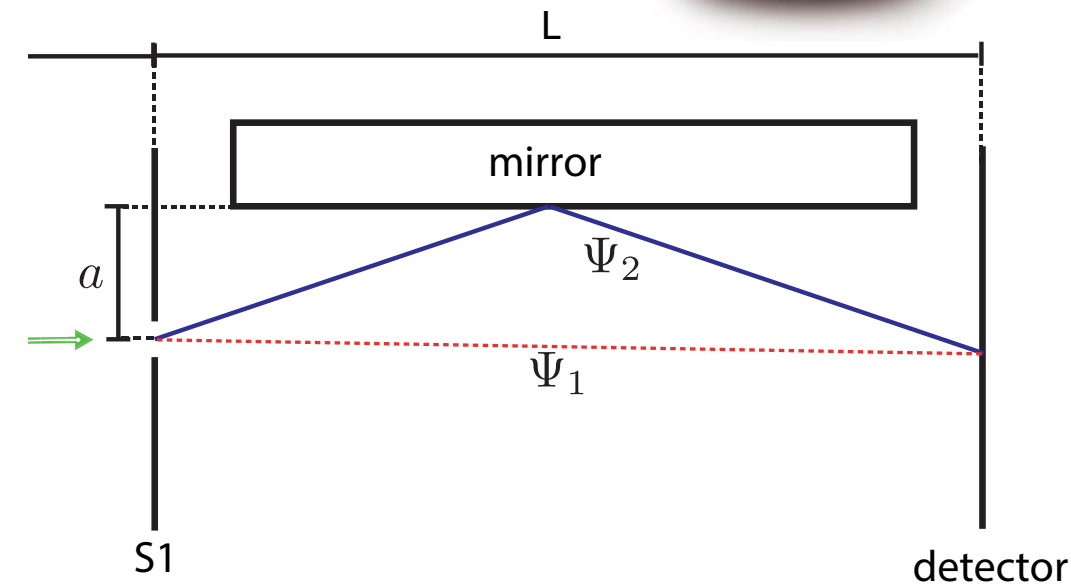
Lloyd interferometer



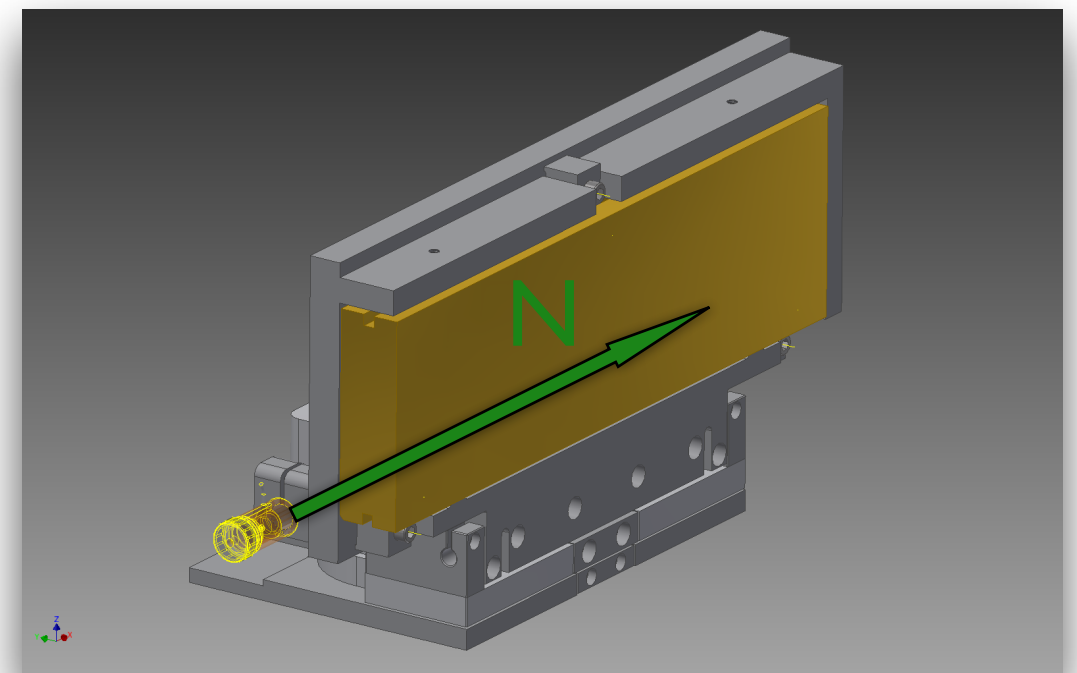
very cold neutrons 10 nm, 40 m/s
in realization by Filter, Masahiro, Oda et al.

$$\lambda_{\text{Lloyd}} = \frac{L}{2a} \lambda_{\text{neutron}}$$

$$\Phi_{\text{Gesamt}} = \Phi_{\text{Geo}} + \Phi_{\text{Refl}} + \Phi_{\text{Grav}} + \Phi_{\text{Corr}} + \Phi_{\text{?}}(y)$$



Simulated MC data



Roman Gergen

Slow neutrons



interferometric setup

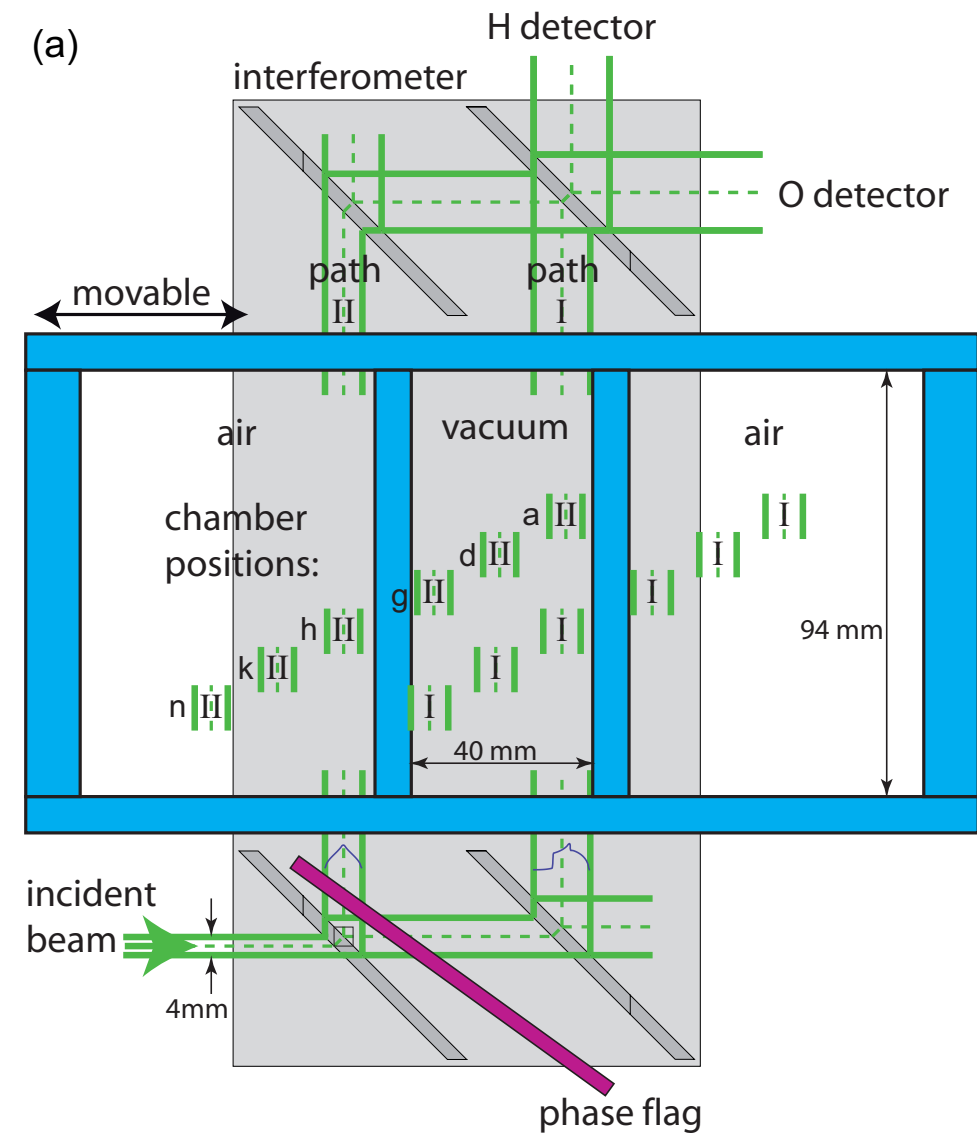
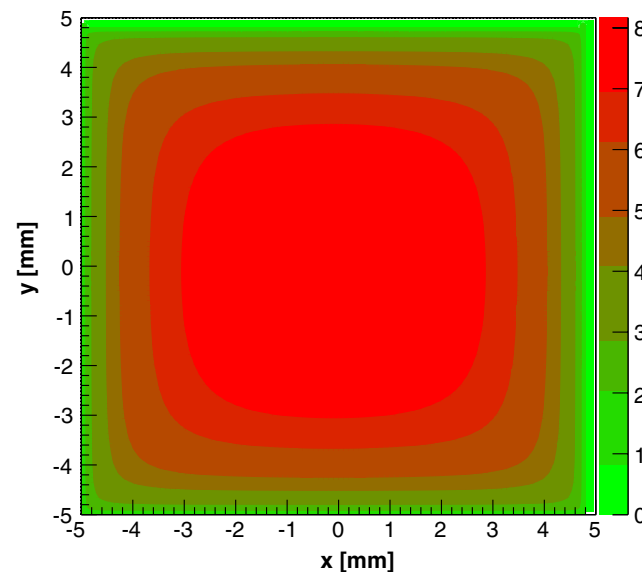
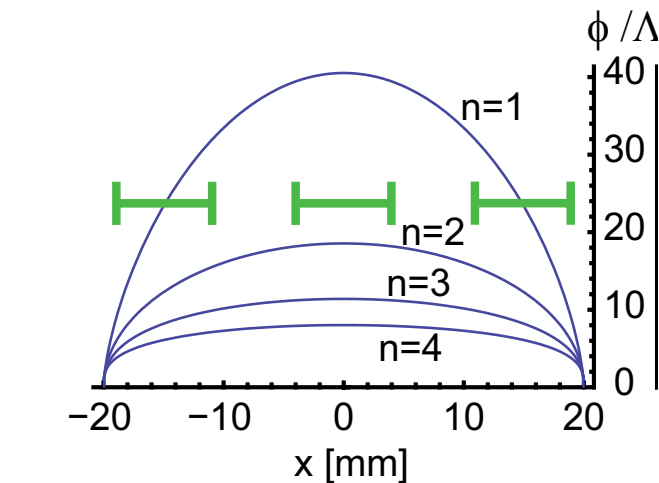
SI 8 at ILL

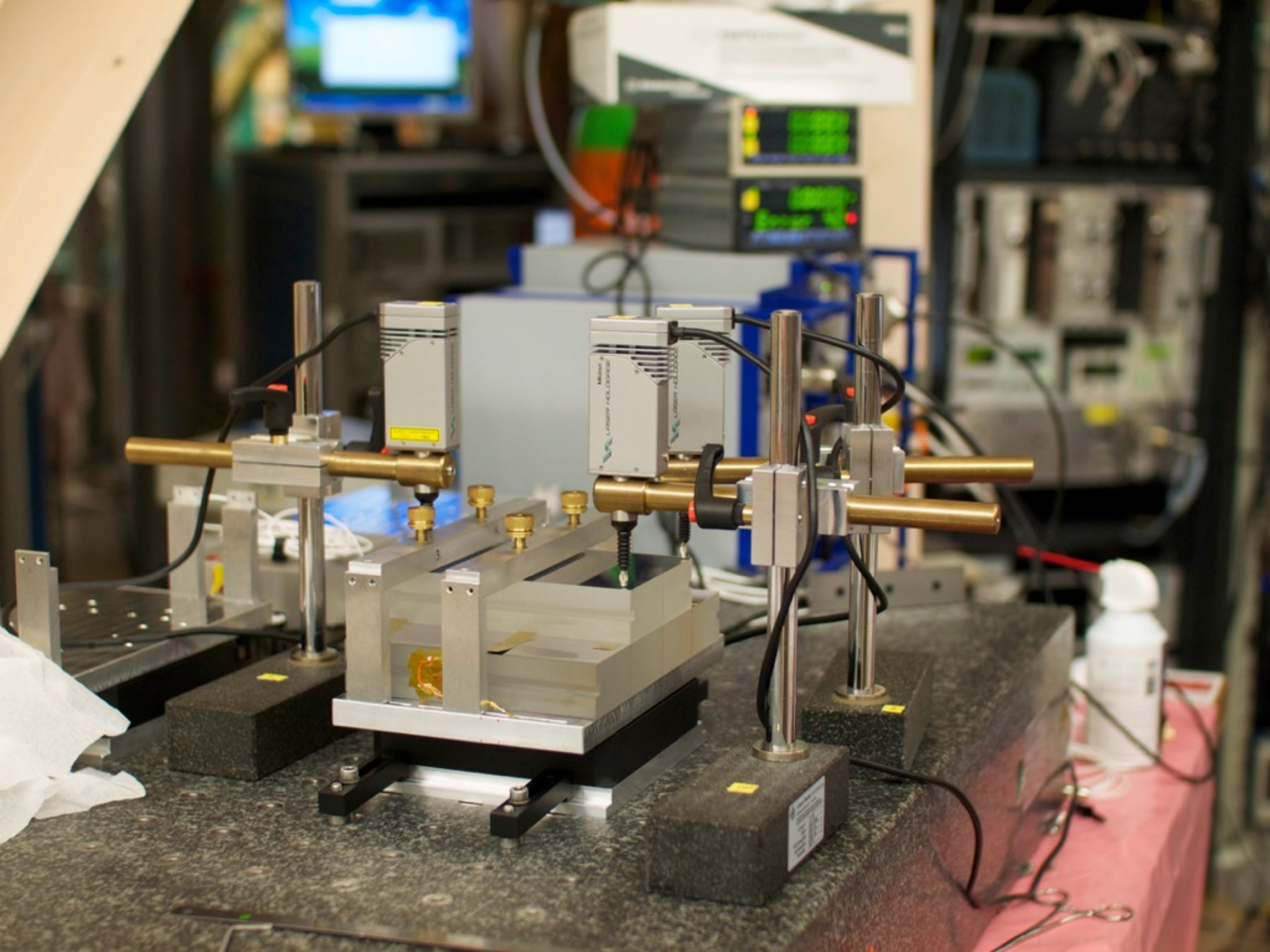
realized by Th. Potocar /
T. Jenke & collaborators

neutrons wavelength

$$\lambda \approx 3\text{\AA}$$

- spacial profile measured
- density dependence







Kevin Mitsch



Jason Jung

Martin Thalhammer



G. C.



Peter Geltenbort



Hanno Filter



Tobias Jenke



Hartmut Abele